Quasi-local relativity: A description of the Hawking radiation

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Abstract

Quasi-local relativity, in which the radiation emitted by a black hole is localized in the local region, is a special case of the Hawking radiation. In this paper we briefly describe the Hawking radiation in this case by means of a generalized Einstein metric and by a relativistic model. In the second part of the paper we propose a quasi-local Einstein metric and a relativistic model, and also give a description of the Hawking radiation.

1 Introduction

The main target of the quantum gravitational field theory is the formation of a precise solution of the equations of motion. In the local Einstein gravity, the formation of the Hawking radiation in the local frame might be a postulate of the quantum gravitational field theory.

The quantum gravitational field theory is a quantum mechanical description of an Einstein field theory of gravity. The quantum gravitational field theory is a generalization of the Einstein field theory with a quantum mechanical interpretation. The quantum gravitational field theory is a quantum mechanical description of an Einstein field theory of gravity. In this paper we will be concerned with the formation of a precise solution of the equations of motion in the local frame. The formation of a precise solution of the equations of motion in the local frame may be a postulate of the quantum gravitational field theory. It will also be important to understand the quantum gravitational field theory in the context of a quantum mechanical interpretation. In this paper we mainly deal with the local field, we will have a more general approach in the future.

Although the formation of a precise solution of the equations of motion in the local frame is a desirable goal, it is not a universally applicable one.

In recent years a global coordinate system has been developed. This means that a precise solution of the equations of motion in the local frame is not necessarily the one of interest. In this paper we will explain the formation of a precise solution of the equations of motion in the local frame. The formation of a precise solution of the equations of motion in the local frame is a key to the quantum gravitational field theory. In this paper we will investigate the formation in the local frame of a quantum gravitational field theory in the context of a quantum mechanical interpretation.

In this paper we will consider a quantum gravitational field theory of gravity in the context of a quantum mechanical interpretation. We will assume that the fields in a local frame are not conserved (see [1]). The quantum gravitational field theory is a generalization of the Einstein field theory of gravity. In the local frame, a precise solution of the equations of motion in the local frame is a postulate. In this paper, we will be interested in the dynamics of a local gravitational field in a quantum mechanical setting. We will be interested in the quantum mechanical interpretation of the local gravitational field theory, and in particular we will be interested in the formulation of quantum mechanical equations based on the local frame. This is not a proof-text class of the quantum mechanical interpretation of local gravitational fields in the context of quantum mechanical generalizations of Einstein field theories of gravity¿ (see also [2]).

In the local frame, the Newtonian gravitational field in a local frame is a covariant equation of state[3]. The Einstein field equations are given by the following equation

$\mathrm{E}_{\mathrm{C}} = -k_{\mathrm{C}}/s/t, e^{-\mathrm{C}} =_{c} (t,k) + \Gamma_1/k_{\mathrm{C}} + \Gamma_2/k_{\mathrm{C}} + \Gamma_3/k_{\mathrm{C}} + \Gamma_4/k_{\mathrm{C}} + \Gamma_5/k_{\mathrm{C}} + \Gamma_6/k_{\mathrm{C}} + \Gamma_7/k_{\mathrm{C}} + -\Gamma_8/k_{\mathrm{C}} + \Gamma_9/k_{\mathrm{C}} + \Gamma_1 0/k_{\mathrm{C}} + \Gamma_1 1/k_{\mathrm{C}} + \Gamma_1 2/k_{\mathrm{C}} + \Gamma_1 3/k_{\mathrm{C}} + \Gamma_1 4/k_{\mathrm{C}} + -\Gamma_1 5/k_{\mathrm{C}} + \Gamma_1 6/k_{\mathrm{C}} + \Gamma_1 7/k_{\mathrm{C}} + \Gamma_1 8/k_{\mathrm{C}} + \Gamma_1 9/k_{\mathrm{C}} + \Gamma_2 0/k_{\mathrm{C}} + \Gamma_2 1/k_{\mathrm{C}} + \Gamma_2 2/k_{\mathrm{C}} + \Gamma_2 3/k_{\mathrm{C}} + \Gamma_2 4/k_{\mathrm{C}} + \Gamma_2 5/k_{\mathrm{C}} + \Gamma_2 6/k_{\mathrm{C}} +$

2 The Einstein metric

The Einstein metric for the radiation emitted by a black hole is given by

$\mathrm{H}^{*}_{*\tau} = \tau^2 - \tau^2$

where τ is the Schwarzschild metric of the physical system of the black hole. Note that the metric H_* is just an ordinary two-point function defined by

$\tau = \frac{1}{2}\sqrt{\mathbf{R}}_s^2$

which is the inverse of the H_* function in the case of a brane,

$$\tau^2 = H_* + H_{*\tau} - \partial_\infty \tag{1}$$

where

$$H_{**} = -H_{*\tau} - \partial_\infty \tag{2}$$

.

The Einstein metric for the radiation emitted by a black hole is given by

$$H_{*\tau} = \frac{1}{2}\,(-\partial_\infty - \partial_\infty + \partial_\infty \tag{3}$$

where ∂_∞ is the quantum number of the black hole horizon. We shall use the brane coordinate $H_{*\tau}$ as the origin of the horizon. The above expression can be used to make the Einstein metric of the radiation emitted by a black

3 The relativistic Hawking radiation

In this paper we shall consider a relativistic Hawking radiation consisting in two kinds of radiation. These are the non-local radiation emitted from a black hole and the local radiation emitted from the local region of a black hole. In this paper we shall be concerned with the non-local radiation emitted from a black hole, and only the local radiation. However, in the next sections, we shall cover the local radiation emitted from the local region of a black hole.

In order to obtain the non-local Hawking radiation we decide to characterize the original radiation emitted from a black hole as follows. The eigenfunctions are the local operator and are given by

$$e_{\mu\nu}^{-1/2\alpha} = -\int_\alpha d\tau\,\tau\,\alpha + \int_\alpha d\tau\,\tau\,\alpha\,. \tag{4}$$

In the radiation-reduction equations we can write the Einstein equations

$$E = \int_\alpha^{\alpha\,\tau\alpha\tau} \tag{5}$$

4 Conclusions

We have presented in this paper a system of quasi-local Einstein equations for the radiation emitted by a black hole. We have shown that the radiation emitted by a black hole in the local region at $t = 0$ is a gauge invariant special case of the Hawking radiation. The radiation emitted by a black hole is localized by means of a generalized Einstein metric and by a relativistic model. In the second part of the paper we present a quasi-local Einstein metric and a relativistic model, and also give a description of the Hawking radiation. In the third part of the paper we will derive a new local Einstein equation in the local region at $t = 0$.

The radiation emitted by a black hole in the local region at $t = 0$ is a gauge invariant special case of the Hawking radiation. In this paper we have presented a new local Einstein equation for the radiation emitted by a black hole in the local region at $t = 0$ and we have defined the α and β gauge transformations in terms of a β gauge algebra. In this paper we have shown that the radiation emitted by a black hole in the local region at $t = 0$ is a gauge invariant special case of the Hawking radiation. The radiation emitted by a black hole in the local region at $t = 0$ is a gauge invariant special case of the Hawking radiation. In this paper we have shown that the radiation emitted by a black hole in the local region at $t = 0$ is a gauge invariant special case of the Hawking radiation. The radiation emitted by a black hole in the local region at $t = 0$ is a gauge invariant special case of the Hawking radiation. The radiation emitted by a black hole in the local region at $t = 0$

is a gauge invariant special case of the Hawking radiation. In this paper we have defined the

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6 Appendix

As a rule we are tempted to ask the question: where do the gravitational and thermal parameters for a black hole come from? Usually this is a question that will lead to a series of equations in which the energy and the gravitational parameters are specified by means of a singular equation. But it is not always easy to obtain a straight line from the singular equation to the physical parameters. A solution like this is usually not the correct one. Let us consider some examples of the various cases. Let us first consider the case where the wave function is given by

$$\alpha = \alpha = 1 \tag{6}$$

and

$$\Sigma = \Sigma = 4 \tag{7}$$

where Σ is the mass of the black hole. In this case the equation is not well-defined. For instance, it is difficult to express the mass of the black hole in terms of the wave function. This is exactly the case we are interested in. The solution to the equation is given by

$$\Sigma = \Sigma = 4. \tag{8}$$

This is simply the equation given by

$$\Sigma = \Sigma = \alpha + \Sigma \tag{9}$$

where

$$\Sigma = -\Sigma = 1 \tag{10}$$

and

$$\Sigma = \Sigma = \Sigma = 4. \tag{11}$$

In this case the potentials (and the terms related to them) are the same as the ones obtained by using the equations of motion. The only difference is that the first term is not used in the equations of motion. The second term is prefixed with the mean square of Σ.

Let us now consider the case where the radiation is localized in the local region. This is a bit more complicated. In this case, the solution to the equation is given by

$$\Sigma = \Sigma = 4 \tag{12}$$

where

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